

Dynamical behavior of interacting dark energy in loop quantum cosmology

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The dynamical behaviors of interacting dark energy in loop quantum cosmology are discussed in this paper. Based on defining three dimensionless variables, we simplify the equations of the fixed points. The fixed points for interacting dark energy can be determined by the Friedmann equation coupled with the dynamical equations in Einstein cosmology. But in loop quantum cosmology, besides the Friedmann equation, the conversation equation also give a constrain on the fixed points. The difference of stability properties for the fixed points in loop quantum cosmology and the ones in Einstein cosmology also have been discussed.

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I. INTRODUCTION

Dark energy [1] is introduced to explain the expanding of our universe. We still have very little knowledge about it. One of the more interested candidate is cosmological constant with the equation of state parameter $\omega_{DE} = -1$. But the physical interpretation of the vacuum energy density is still faultiness. And the problem that why the dark energy and dark matter are at the same order today is still unknown, this is so-called coincidence problem. To overcome this problem, many dynamical dark energy models have been introduced to replace the cosmological constant. But there are still many other approaches to explain the coincidence problem. An interesting proposal is that the dark energy is interacting with dark matter. Many interacting models have been studied. The interacting term always is the function of the Hubble parameter H , the density of dark energy ρ_{DE} and the density of the dark matter ρ_{DM} (for example, please see the second section of [2] and the references of it), or the function of Γ_x , ρ_{DE} and ρ_{DM} [3–7].

A general interaction between dark energy and dark matter is described by

$$\dot{\rho}_{DE} + 3H\rho_{DE}(1 + \omega_{DE}) = Q, \quad (1)$$

$$\dot{\rho}_{DM} + 3H\rho_{DM} = -Q. \quad (2)$$

in which ω_{DE} is the equation of state parameter of dark energy. A constant ω_{DE} and $\omega_{DE} = -1$ crossing were studied by many authors [7–9]. In this paper we just consider ω_{DE} is a negative constant. Q is the interacting term. And we will discuss two models:

$$\text{Model I} \quad Q_1 = 3HC_1\rho_{DE} + 3HC_2\rho_{DM}, \quad (3)$$

$$\text{Model II} \quad Q_2 = 3^{2-n}CH^{3-2n}\rho_{DM}^n, \quad (4)$$

in which C_1, C_2, C are some constants. Model I has been discussed by [2, 7] in Einstein cosmology (EC), they considered that $C_1 \neq C_2$. But more special cases also have

been studied, e.g., $C_1 = C_2$ [10], $C_1 = 0$ [11], $C_2 = 0$ [12], in EC. For more simplicity, we will discuss $C_1 = C_2 = C$ for the first model. The second model also has been discussed in EC [13, 14], n is constant, and we will discuss $n = 2$. We will introduce three dimensionless variables to study those models in loop quantum cosmology (LQC), and discuss the difference and the common properties for dynamical system between LQC and in EC.

LQC [15–17] is a canonical quantization of homogeneous spacetimes based upon techniques used in loop quantum gravity. Due to the homogeneous and isotropic spacetime, the phase space of LQC is simpler than loop quantum gravity, e.g., the connection is determined by a single parameter called c and the triad is determined by p . The variables c and p are canonically conjugate with Poisson bracket $\{c, p\} = \frac{\gamma}{3}$, in which γ is the Barbero-Immirzi parameter. In LQC scenario, the initial singularity is instead of a bounce. So thanks for the quantum effect, the universe is initially contracting phase with the minimal but not zero volume, and then the quantum effect drives it to the expanding phase. In the effective scenario, the behavior of LQC also can be well described, e.g., the big crunch instead of the big bounce [18, 19]. People always consider two kinds of correction of effective LQC, one is inverse volume correction, the other one is holonomy correction. In this paper, we just discuss the holonomy correction. In this effective LQC, the Friedmann equation is modified (see Eqs.(24)). It is easy to get the difference between the Friedmann equation in EC and the one in LQC: it adds a $-\frac{\rho^2}{\rho_c}$ term to standard Friedmann equation which essentially encodes the discrete quantum geometric nature of spacetime [20, 21]. This means that the energy density will have a critical value, when the energy density is very nearly the critical density ρ_c , the cosmology will have a bounce and then oscillates forever. Due to the modification of the Friedmann equation, the dynamical behaviors of dark energy and dark matter are very different from the ones in EC. This factor was showed by many authors, for example, the phantom field [22], the quintom field [23], the interacting phantom field [24], the interacting dark energy with the interacting term $Q = 3cH\rho_{DE}^\alpha\rho_{DM}^{1-\alpha}$ [11], and the dark energy interacting with cold dark matter with-

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out coupling to the baryonic matter[25]. In this paper, we will discuss the dynamical behavior of interacting dark energy.

This paper is organized as following. In Sec. II, we will review the dynamical behavior for interacting dark energy in EC. And Sec. III will show the dynamical behavior of interacting dark energy in LQC. The numerical analysis will be showed in the Sec. IV. And we get some conclusions in the last Sec. V. For simplicity, we set $8\pi G = 1$.

II. DYNAMICAL BEHAVIOR IN EC

In EC, the Friedmann equation can be written as

$$H^2 = \frac{1}{3}\rho, \quad (5)$$

in which $\rho = \rho_{DE} + \rho_{DM}$ is the total energy density. Differentiating Eq.(5) and using the conversation law of the total energy density $\dot{\rho}_t + 3H(\rho + p) = 0$, we can get

$$\dot{H} = -\frac{1}{2}(\rho + p), \quad (6)$$

in which $p = p_{DE}$ is the pressure of dark energy (the pressure of dark matter is zero), and $p_{DE} = \rho_{DE}\omega_{DE}$, with the equation of state parameter ω_{DE} . We just consider ω_{DE} is a negative constant, then it is another parameter. To discuss the dynamical behavior of interacting dark energy in EC, we introduce the dimensionless variables [7, 26]:

$$x = \frac{\rho_{DE}}{3H^2}, \quad y = \frac{\rho_{DM}}{3H^2}, \quad z = \frac{Q}{3H^3}. \quad (7)$$

Although the density of dark energy can be a negative on in the f(R) theory [27], we restrict that the density of dark energy and dark matter both should be positive or one of them is zero. So both of x and y should be non-negative.

The Friedmann equation (5) and Eq.(6) can be rewritten as

$$x + y = 1, \quad (8)$$

$$\frac{\dot{H}}{H^2} = -\frac{3}{2} \left(1 + \frac{\omega_{DE}x}{x+y} \right) = -\frac{3}{2} (1 + \omega_{DE}x). \quad (9)$$

The effective total equation of state parameter ω is

$$\omega = \frac{\rho_{DE}\omega_{DE}}{\rho_{DE} + \rho_{DM}} = \frac{\omega_{DE}x}{x+y} = \omega_{DE}x. \quad (10)$$

In EC, the dynamical behavior of dimensionless variable x, y are

$$x' = z - 3\omega_{DE}x + 3\omega_{DE}x^2, \quad (11)$$

$$y' = -z + 3\omega_{DE}xy. \quad (12)$$

in which prime denotes the derivative with respect to the e-folding number $N = \ln a$ ($a_0 = 1$). Just as [7]

argued, if $z = z(x, y)$, the above equations are closed and autonomous. The model I (3) and II (4) both are in this case.

For a autonomous system $\mathbf{X}' = f(\mathbf{X})$, the fixed points \mathbf{X}_* satisfy $\mathbf{X}' = 0$. So, setting $x' = 0, y' = 0$ of Eqs.(11,12) and simplifying them, it is easy to get the fixed points satisfy:

$$x_* + y_* = 1, \quad (13a)$$

$$-3\omega_{DE}x_*(1 - x_*) + z_* = 0, \quad (13b)$$

in which z_* is the value of z in the fixed points. If $x_* = 0$, the second subequation of Eq.(13) implies that $z_* = 0$. This is a point for pure dark matter dominated [7], we will not consider this condition. In this paper, we will discuss two interacting model: $Q_1 = 3HC_1\rho_{DE} + 3HC_2\rho_{DM}$, and $Q_2 = 3^{2-n}H^{3-2n}C\rho_{DM}^n$. C_1, C_2 and C are some constants. For simplicity, we will consider $C_1 = C_2 = C$ and $n = 2$ for those two models. Those models have been studied in EC by many authors [2, 7–14]. But for convenience to compare the dynamical behavior of interacting dark energy in LQC and the ones in EC, we review the dynamical behaviors of dark energy with different interacting term in EC in the next two subsections.

A. Model I: $Q = 3HC_1\rho_{DE} + 3HC_2\rho_{DM}$

This interacting term has been discussed in many papers [2, 7–12]. In this subsection, for more simplicity, we just consider $C_1 = C_2 = C$, the dimensionless interacting term is $z = 3C(x + y)$. Eq.(13) can be rewritten as

$$-\omega_{DE}x_*(1 - x_*) + C = 0, \quad (14)$$

It is easy to get the fixed points (x_*, y_*) from Eq.(14):

$$\text{Point } A_1 : \left(\frac{1 + \sqrt{1 - 4\frac{C}{\omega_{DE}}}}{2}, \frac{1 - \sqrt{1 - 4\frac{C}{\omega_{DE}}}}{2} \right) \quad (15)$$

$$\text{Point } B_1 : \left(\frac{1 - \sqrt{1 - 4\frac{C}{\omega_{DE}}}}{2}, \frac{1 + \sqrt{1 - 4\frac{C}{\omega_{DE}}}}{2} \right) \quad (16)$$

In order to determine the stability properties of the fixed points, it is necessary to expand the autonomous system $\mathbf{X}' = f(\mathbf{X})$ around the fixed points \mathbf{X}_* , setting $\mathbf{X} = \mathbf{X}_* + \mathbf{U}$ with \mathbf{U} the perturbations of the variables considered as a column vector. Thus, for each fixed point, one can expand the equation for the perturbations up to the first order as $\mathbf{U}' = \mathbf{M} \cdot \mathbf{U}$, where the matrix \mathbf{M} contains the coefficients of the perturbation equations. The stability property for each fixed point is determined by the eigenvalues of \mathbf{M} [14]. According to this property, it

is easy to get the eigenvalues for the autonomous system:

$$\begin{aligned} \text{Point } A_1 : \quad \lambda_1 &= \frac{3}{2}\sqrt{\omega_{DE}(\omega_{DE} - 4C)} + \frac{3}{2}\omega_{DE}, \\ \lambda_2 &= 3\sqrt{\omega_{DE}(\omega_{DE} - 4C)}, \end{aligned} \quad (17)$$

$$\begin{aligned} \text{Point } B_1 : \quad \lambda_1 &= -\frac{3}{2}\sqrt{\omega_{DE}(\omega_{DE} - 4C)} + \frac{3}{2}\omega_{DE}, \\ \lambda_2 &= -3\sqrt{\omega_{DE}(\omega_{DE} - 4C)}. \end{aligned} \quad (18)$$

The stability property for each fixed point is determined by the sign of λ_1, λ_2 . The fixed point is a stable if both of λ_1, λ_2 are negative, and unstable if both of them are negative, if λ_1, λ_2 have not the same sign, then the fixed point is a saddle point. To ensure that $0 < x_*, y_* < 1$, $0 < 4\frac{C}{\omega_{DE}} < 1$ should be satisfied. For Point A_1 , it is always a saddle point for $\lambda_1 < 0$ and $\lambda_2 > 0$. For Point B_1 , it is a stable point always satisfied, for both of λ_1, λ_2 always are negative if $0 < 4\frac{C}{\omega_{DE}} < 1$ is held. If total equation of state parameter $\omega = \frac{\omega_{DE} - \omega_{DE}\sqrt{1-4\frac{C}{\omega_{DE}}}}{2}$ is in the regions of $-1 < \omega < -\frac{1}{3}$, the universe will enter an accelerated stage, then Point B_1 is an accelerated scaling attractor, this has been shown in the [2]. If $\omega < -1$, it is easy to get $\dot{H} > 0$ from Eq.(9), the universe will have a big rip.

B. Model II: $Q = 3^{2-n}H^{3-2n}C\rho_{DM}^n$

This interacting term has been discussed by [13, 14]. In this section, we just consider $n = 2$, the dimensionless variable is $z = 3Cy^2$, then Eqs.(13) can be rewritten as

$$-\omega_{DE}x_*(1-x_*) + C(1-x_*)^2 = 0. \quad (19)$$

The fixed points (x_*, y_*) for this model are

$$\text{Point } A_2 : \quad (1, 0), \quad (20)$$

$$\text{Point } B_2 : \quad \left(\frac{C}{\omega_{DE} + C}, \frac{\omega_{DE}}{\omega_{DE} + C} \right). \quad (21)$$

And the eigenvalues are

$$\text{Point } A_2 : \quad \lambda_1 = 3\omega_{DE}, \quad \lambda_2 = 3\omega_{DE}, \quad (22)$$

$$\text{Point } B_2 : \quad \lambda_1 = -3\omega_{DE}, \quad \lambda_2 = \frac{3\omega_{DE}C}{C + \omega_{DE}}. \quad (23)$$

For Point A_2 , the condition of $x_*, y_* > 0$ and $\omega_{DE} < 0$ constrain that $C < 0$ and $\omega_{DE} + C < 0$ should be held. It is always a stable point for $\omega_{DE} < 0$. The universe will enter a dark energy dominate stage, and $\omega = \omega_{DE}$. For Point B_2 , $C < 0$ should be satisfied, it is always a saddle point. So, for this interacting model, it has not any attractive behavior except the universe enter a dark energy dominated stage.

III. DYNAMICAL BEHAVIOR IN LQC

In this section, we discuss the dynamical behavior of interacting dark energy in LQC. As mentioned in the

introduction, the Friedmann equation is modified in LQC and can be written as

$$H^2 = \frac{1}{3}\rho \left(1 - \frac{\rho}{\rho_c} \right), \quad (24)$$

in which $\rho_c = \frac{3}{\kappa\gamma^2\alpha\ell_{Pl}^2}$. Differentiating Eq.(24) and using the conversation law of the total energy density $\dot{\rho}_t + 3H(\rho + p) = 0$, one can get

$$\dot{H} = -\frac{1}{2}(\rho + p) \left(1 - \frac{2\rho}{\rho_c} \right), \quad (25)$$

Considering the dimensionless variables (7), we can rewritten Eqs.(24,25) as

$$(x + y) \left[1 - \frac{3H^2}{\rho_c}(x + y) \right] = 1, \quad (26)$$

$$\frac{\dot{H}}{H^2} = -\frac{3}{2} \left[1 + \frac{x\omega_{DE}}{x + y} \right] (2 - x - y). \quad (27)$$

The total effective equation of state parameter ω is

$$\omega = \frac{\rho_{DE}\omega_{DE}}{\rho_{DE} + \rho_{DM}} = \frac{\omega_{DE}x}{x + y}. \quad (28)$$

Considering the dimensionless variables, the dynamical equations for our interesting system can be obtained:

$$\begin{aligned} x' &= z - 3x(1 + \omega_{DE}) \\ &+ 3x \left(1 + \frac{x\omega_{DE}}{x + y} \right) (2 - x - y), \end{aligned} \quad (29)$$

$$y' = -z - 3y + 3y \left(1 + \frac{x\omega_{DE}}{x + y} \right) (2 - x - y), \quad (30)$$

in which prime denotes the derivative with respect to the e-folding number $N = \ln a$ ($a_0 = 1$).

The fixed points (x_*, y_*) satisfy $x' = 0, y' = 0$. Simplifying Eqs.(29,30), it is easy to get

$$\text{Case I} \quad x_* + y_* = 1, \quad (31a)$$

$$-3\omega_{DE}x_*(1 - x_*) + z_* = 0. \quad (31b)$$

or

$$\text{Case II} \quad x_*(1 + \omega_{DE}) + y_* = 0, \quad (32a)$$

$$-3x_*(1 + \omega_{DE}) + z_* = 0. \quad (32b)$$

Comparing the above equations with Eqs.(13), we find that the expression for the first case is as same as in the EC. The Case II is a special case which just exists in LQC. This is because of the Friedmann equation has been modified, then the dynamical equations in LQC are different from the ones in EC (see Eqs.(29,30) and Eqs.(11,12)). We will discuss these two cases for fixed point in the last section.

As in EC, the values of fixed points depend on the special form of dimensionless interacting term z . We will consider two models: $Q = 3HC_1\rho_{DE} + 3HC_1\rho_{DM}$, and $Q = 3^{2-n}H^{3-2n}C\rho_{DM}^n$, with some constants C_1, C_2, C, n .

A. Model I: $Q = 3HC_1\rho_{DE} + 3HC_2\rho_{DM}$

In this model, the dimensionless variable z can be written as

$$z = \frac{Q}{3H^3} = 3C_1x + 3C_2y, \quad (33)$$

in which C_1, C_2 are some constants.

Considering the above equation, Case I and Case II can be rewritten as

$$\omega_{DE}x_*^2 - \omega_{DE}x_* + C_1x_* + C_2(1 - x_*) = 0, \quad (34)$$

$$C_1x_* + C_2[-x_*(1 + \omega_{DE})] - x_*(1 + \omega_{DE}) = 0 \quad (35)$$

The first equation corresponds to Case I, and the second one corresponds to Case II. It is easy to find that $x_* = 0$ in the Case II. This means that $y_* = 0$. The universe will enter other stuff dominated stage, this is beyond our interesting, so we will ignore this case in this subsection. As the last section, we just consider $C_1 = C_2 = C$.

Now, we consider Case I. Eq.(34) becomes

$$x_*^2 - x_* + \frac{C}{\omega_{DE}} = 0. \quad (36)$$

Then the fixed points (x_*, y_*) are

$$\text{Point } A_3 : \left(\frac{1 + \sqrt{1 - 4\frac{C}{\omega_{DE}}}}{2}, \frac{1 - \sqrt{1 - 4\frac{C}{\omega_{DE}}}}{2} \right) \quad (37)$$

$$\text{Point } B_3 : \left(\frac{1 - \sqrt{1 - 4\frac{C}{\omega_{DE}}}}{2}, \frac{1 + \sqrt{1 - 4\frac{C}{\omega_{DE}}}}{2} \right) \quad (38)$$

The eigenvalues are

$$\begin{aligned} \text{Point } A_3 : \quad \lambda_1 &= -3 - \frac{3}{2}\omega_{DE} - \frac{3}{2}\sqrt{\omega_{DE}(\omega_{DE} - 4C)}, \\ \lambda_2 &= 3\sqrt{\omega_{DE}(\omega_{DE} - 4C)}, \end{aligned} \quad (39)$$

$$\begin{aligned} \text{Point } B_3 : \quad \lambda_1 &= -3 - \frac{3}{2}\omega_{DE} + \frac{3}{2}\sqrt{\omega_{DE}(\omega_{DE} - 4C)}, \\ \lambda_2 &= -3\sqrt{\omega_{DE}(\omega_{DE} - 4C)}. \end{aligned} \quad (40)$$

The density of dark energy and dark matter both are positive ones or one of them is zero, this condition given a constraint on the relationship between ω_{DE} and C . In this case, $\omega_{DE} < 0$ and $0 < \frac{C}{\omega_{DE}} \leq \frac{1}{4}$ both should be satisfied. For Point A_3 , it is a unstable point if $1 + \omega_{DE} + \omega_{DE}C > 0$ and $\omega_{DE} < -2$, and it is a saddle point if $\omega_{DE} > -2$ or $1 + \omega_{DE} + \omega_{DE}C < 0$ and $\omega_{DE} < -2$. For Point B_3 , it is stable point if $\omega_{DE} > -2$ and $1 + \omega_{DE} + \omega_{DE}C > 0$, and a saddle point if $\omega_{DE} > -2$ and $1 + \omega_{DE} + \omega_{DE}C < 0$ or $\omega_{DE} < -2$.

Comparing the above fixed points and eigenvalues with the ones of classical scenario, we find the expressions for fixed points are same for it is not necessary to consider Case II, but the eigenvalues are different. The reason

is that the eigenvalues are determined by the dynamical equations, the dynamical equations in LQC is different from the ones in EC. In EC, the singularity maybe happens. But in LQC, for Friedmann equation has been modified, the singularity is replaced by the bounce always. We will show this evidence in the next section.

B. Model II: $Q = 3^{2-n}CH^{3-2n}\rho_{DM}^n$

In the model, the dimensionless interacting term is $z = 3Cy^n$, then the Eqs.(31,32) can be rewritten as

$$\omega_{DE}x_*^2 - \omega_{DE}x_* + C(1 - x_*)^n = 0, \quad (41)$$

$$(-1)^nC(1 + \omega_{DE})^nx_*^n - x_*(1 + \omega_{DE}) = 0. \quad (42)$$

To get the values of x_* , we need the special value of n . As an example to use Case II to get the fixed points, we will consider $n = 2$. Then, the dimensionless variable z is $z = 3Cy^2$. In this case, Eqs.(41,42) can be rewritten as

$$\omega_{DE}x_*^2 - \omega_{DE}x_* + C(1 - x_*)^2 = 0, \quad (43a)$$

$$Cx_*^2(1 + \omega_{DE})^2 - x_*(1 + \omega_{DE}) = 0. \quad (43b)$$

Solving the above equations, it is easy to get the fixed points

$$\text{Point } A_4 : (1, 0), \quad (44)$$

$$\text{Point } B_4 : \left(\frac{C}{\omega_{DE} + C}, \frac{\omega_{DE}}{\omega_{DE} + C} \right), \quad (45)$$

$$\text{Point } C_4 : \left(\frac{1}{C(1 + \omega_{DE})}, -\frac{1}{C} \right). \quad (46)$$

Noticed that $x_* = 0$ is also the solution of Eq.(43b), but considering the first equation of Eq.(32), y_* should be zero when $x_* = 0$, the universe will enter other stuff dominated, this is beyond our interesting.

The eigenvalues for above points are

$$\text{Point } A_4 : \quad \lambda_1 = -3(1 + \omega_{DE}), \quad \lambda_2 = 3\omega_{DE}, \quad (47)$$

$$\begin{aligned} \text{Point } B_4 : \quad \lambda_1 &= -3\omega_{DE}, \\ \lambda_2 &= -3\frac{C + C\omega_{DE} + \omega_{DE}}{\omega_{DE} + C}, \end{aligned} \quad (48)$$

$$\begin{aligned} \text{Point } C_4 : \quad \lambda_1 &= -\frac{3}{2}\omega_{DE} \\ &- \frac{3}{2}\frac{\sqrt{\omega_{DE}^2C^2 - 4C^2 - 4\omega_{DE}C^2 - 4C\omega_{DE}}}{C}, \\ \lambda_2 &= -\frac{3}{2}\omega_{DE} \\ &+ \frac{3}{2}\frac{\sqrt{\omega_{DE}^2C^2 - 4C^2 - 4\omega_{DE}C^2 - 4C\omega_{DE}}}{C}. \end{aligned} \quad (49)$$

For Point A_4 , the universe will enter a dark energy dominated stage. If $-1 < \omega_{DE} < -\frac{1}{3}$, e.g., the dark energy is quintessence-like type, this point is a stable

point, and it is an accelerated scaling attractor. If $\omega_{DE} < -1$, it is a saddle point. For Point B_4 , to ensure $0 < x_*, y_* < 1$, $C < 0$ should be satisfied. So it is a saddle point if $C + C\omega_{DE} + \omega_{DE} < 0$ and an unstable point if $C + C\omega_{DE} + \omega_{DE} > 0$. For Point C_4 , $C < 0, \omega_{DE} < -1$ and $\omega_{DE}^2 C - 4C - 4\omega_{DE} C - 4\omega_{DE} < 0$ need to be satisfied, it is a unstable point if $C + C\omega_{DE} + \omega_{DE} < 0$ and a saddle point if $C + C\omega_{DE} + \omega_{DE} > 0$.

So, this model in LQC will have an attractor behavior just when the universe enter a dark energy with $-1 < \omega_{DE} < -\frac{1}{3}$ dominated stage. This is different from the one of EC, in which the dynamical system will have an attractor behavior when the universe is all types of dark energy dominated. As the first model, due to the quantum modification of Friedmann equation, the dynamical behavior for LQC is different from the one for EC.

IV. NUMERICAL ANALYSIS

In this section, we will analysis the dynamical behavior of interacting dark energy by using numerical tool.

In this paper, we restricted that $0 \leq x_*, y_* \leq 1$. This restriction gives a constraint on C, ω_{DE} . As discussion before, although the expressions for the fixed points in LQC are as same as the ones in EC for Model I, but due to the different dynamical equations, the stability behaviors for each fixed points in LQC are different from the ones in EC. We show the stable regions in the parameter space (ω_{DE}, C) for the Model I in Fig.1. It is easy to find that the stable region for LQC is smaller then the one for EC.

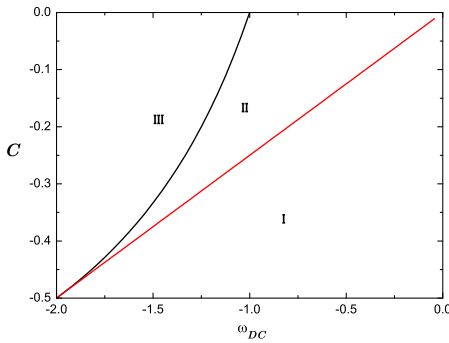


FIG. 1: The stable regions in the (ω_{DE}, C) parameter space for Model I. In EC, the stable regions are region II + III. In LQC, the stable region is region II.

In Fig.2, we plot the evolution of the total equation of state parameter ω_t (described by Eqs.(10,28)) for different models. As Eqs.(10,28) showed, the value of $\omega = \frac{\rho_{DE}\omega_{DE}}{\rho_{DE} + \rho_{DM}}$ is determined by the values of ρ_{DE}, ρ_{DM} , and also depends on the interacting term. But just as the diagram showed, the final value of ω will tend to constant. When $\omega_{DE} = -0.6, C = -0.1$, we can get the

total effective equation of state parameter $-1 < \omega < -\frac{1}{3}$ in the final data for Model I, not only in EC, but also in LQC. But for model II, ω will be less then -1 , and the universe will be phantom-like dark energy dominated. In EC, the universe will have a singularity, but in LQC, the universe will enter an oscillatory region and will have not any singularity, just as Fig.3 shows.

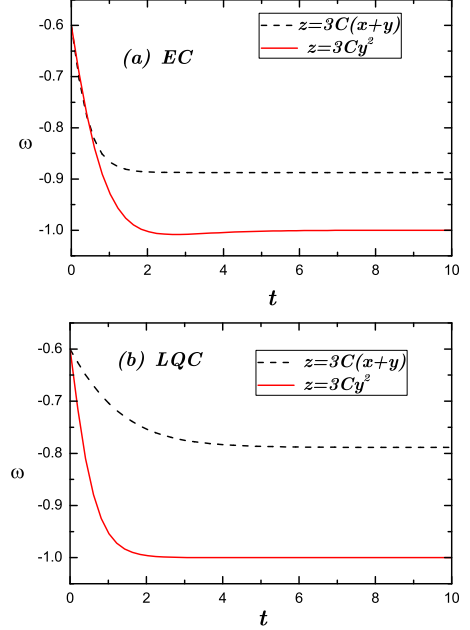


FIG. 2: (Color online). The evolution ω in LQC. The values of constants are $C = -0.1, \omega_{DE} = -0.6, \rho_c = 0.82$. The upper panel is for EC, and the under one is for LQC.

Figure 4 shows the stable point in the classical cosmology and in LQC for Model I. Although the expressions of the fixed points in the classical cosmology are as same as the ones in LQC, but due to the dynamical equations are different, so the stable region in (ω_{DE}, C) space is very different, just as Fig.(1) shows, then the value of stable points for those two cosmology are different.

In Fig.5, we show the evolution of the densities in EC scenario and LQC scenario. In EC, the densities will become bigger and bigger. But in LQC scenario, the densities of dark energy and dark matter and the total matter will not divergence. The universe will entre a oscillating regime, and the singularity is avoided.

V. DISCUSSION AND CONCLUSION

In this paper, based on defining some dimensionless variables, we simplified the equations of fixed points. We found $x_* + y_* = 1$ for the fixed points should be satisfy both in EC and LQC. It is easy to understand this condition hold in EC, for $x + y = 1$ is nothing but the Friedmann equation, just as Eq.(8) shows. But in LQC,

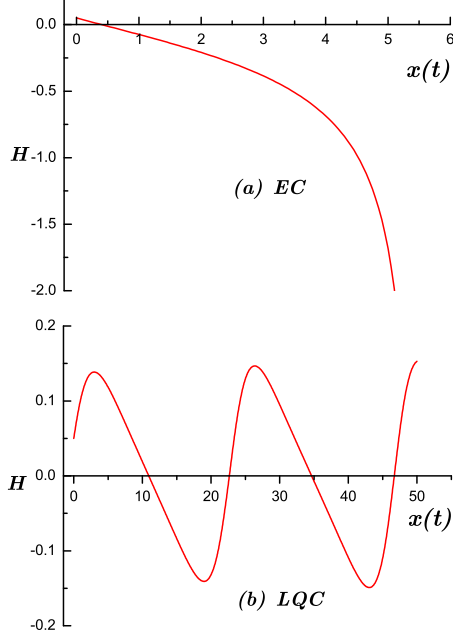


FIG. 3: The evolution of Hubble parameter in the Model II. The values of constants are $C = -0.05$, $\omega_{DE} = -0.6$, $\rho_c = 0.82$. The upper panel is for EC, and the under one is for LQC.

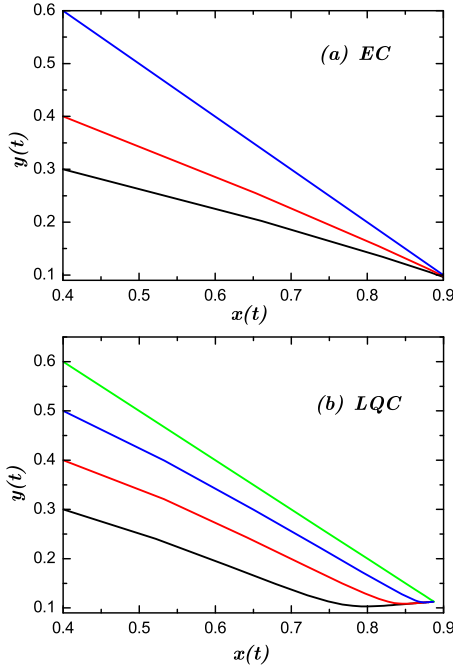


FIG. 4: (Color online). The phase diagrams of interacting dark energy. The upper panel is for EC, with $\omega_{DE} = -1.2$, $C = -0.1$. And the under one is for LQC, with $\omega_{DE} = -1$, $C = -0.1$, $\rho_c = 0.82$.

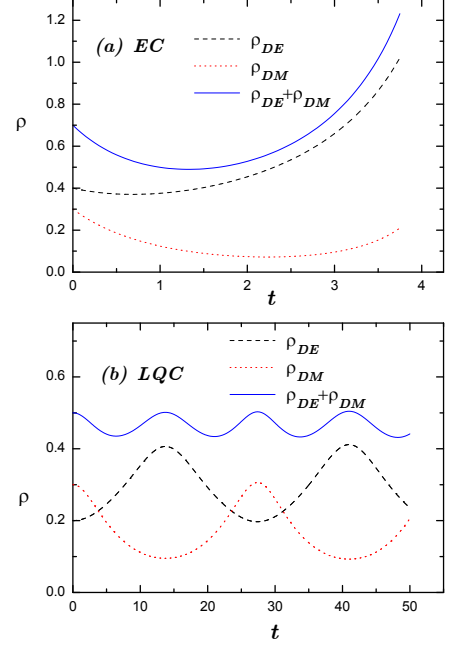


FIG. 5: (Color online). The evolution of energy density ρ . The upper one is the evolutions of ρ_{DE} , ρ_{DM} and the total energy density $\rho = \rho_{DE} + \rho_{DM}$ in EC. The under one is for those variables in LQC. Both of those two diagrams are for $\omega_{DE} = -1.6$, $C = -0.1$.

if the quantum effect is dominated, it is impossible to be satisfied for $\rho < \rho_c$. But notice that the quantum effect should be considered just when ρ is comparable to ρ_c . So $x_* + y_* = 1$ means that our universe is described by general relativity very well when the dynamical system arrives its fixed point. The energy density is very far from the one should consider the quantum effect. It is easy to check that all fixed points of the interacting models discussed in [11] are determined by the condition of Case I. Case II is just held in LQC, but we can get this case from the conversation equation for energy density

$$\dot{\rho} = -3H[\rho_{DE}(1 + \omega_{DE}) + \rho_{DM}]. \quad (50)$$

When $\dot{\rho}_{DE} = \dot{\rho}_{DM} = 0$, this implies $\rho_{DE}(1 + \omega_{DE}) + \rho_{DM} = 0$ and $x' = y' = 0$. Although the conversation for energy density should be satisfied both in EC and LQC, $x_*(1 + \omega_{DE}) + y_* = 0$ means that $x_*\omega_{DE} = -1$ in classical universe, but considering Eqs.(11,12) and setting $x' = y' = 0$, this is nothing but $x_* + y_* = 1$. So, Case II is just a special condition of Case I in the region where the universe is described by general relativity. When the universe is quantum effect dominated, $x_* + y_* = 1$ will not be satisfied, so, Case II is a special condition in LQC. Considering Eq.(32), we can find that the fixed points are described by this case is happening when $z_* \propto x_*^n$ with $n \neq 0, 1$. Notice that, $\omega_{DE} < -1$ should be satisfied in this case for $\rho_{DE}, \rho_{DM} > 0$. Those fixed points which

are determined by Case II are staying in the quantum regions, and if they are stable points, the universe will not enter the Einstein cosmology.

Due to the Friedmann equation is modified in LQC, the dynamical equations in LQC are different from the ones in EC. Under the help of the simplified equations of fixed points, we discussed the dynamical behavior for two different interacting models in LQC. We find that the stability of system dependent on the parameters ω_{DE}, C . Also, $0 \leq x_*, y_* \leq 1$ gives a constraint on ω_{DE}, C . The attractor solutions for autonomous system in LQC are very different from the ones in EC, this fact is obvious when we compare the stable regions in (ω_{DE}, C) spaces of two conditions, although the attractor solution is lying in the stage which is described by general relativity

both for LQC and EC. The values of those parameters also should be constrained by the observation data, it is discuss by many authors [30–33] in EC, but it is still need more studying in LQC. Also, just as we expecting, the unstable region in LQC, the universe will enter an oscillatory regime and avoid the singularity which always happens in EC.

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- [1] Edmund J. Copeland, M. Sami and Shinji Tsujikawa, Int. J. Mod. Phys. D **15**, 1753(2006).
 - R. Rakhi and K. Indulekha, *Dark Energy and Tracker Solution: A Review*, arXiv:0910.5406.
 - Aruna Kesavan, *Dark energy*, arXiv:0908.2852.
 - [2] Miguel Quartin, Mauricio O. Calvao, Sergio E. Joras, Ribamar R. R. Reis and Ioav Waga, JCAP **0805**, 007(2008).
 - [3] M.S. Turner, Phys. Rev. D **28**, 1243(1983).
 - [4] K.A. Malik, D. Wands and C. Ungarelli, Phys. Rev. D **67**, 063516(2003).
 - [5] M. R. Setare and E. C. Vagenas, Int. J. Mod. Phys. D **18**, 147(2009).
 - [6] M. R. Setare and E. C. Vagenas, Phys. Lett. B **666**, 111(2008).
 - [7] G. Caldera-Cabral, R. Maartens and L. A. Urena-Lopez, Phys. Rev. D **79**, 063518(2009).
 - [8] H. Mohseni Sadjadi and M. Alimohammadi, Phys. Rev. D **74**, 103007(2006).
 - [9] Sergio del Campo, Ramon Herrera and Diego Pavon, Phys. Rev. D **78**, 021302(R)(2008).
 - [10] Luis P. Chimento, Alejandro S. Jakubi, Diego Pavon and Winfried Zimdahl, Phys. Rev. D **67**, 083513(2003).
 - [11] S.B. Chen, B. Wang, J.L. Jing, Phys. Rev. D **78**, 123503(2008).
 - [12] Andrew P. Billyard and Alan A. Coley, Phys. Rev. D **61**, 083503(2000).
 - [13] Xi-ming Chen and Yungui Gong, Phys. Lett. B **675**, 9(2009).
 - [14] Xi-ming Chen Yungui Gong and Emmanuel N. Saridakis, JCAP **0904**, 001(2009).
 - [15] Martin Bojowald, Living Rev. Rel. **11**, 4(2008).
 - [16] Abhay Ashtekar, J. Phys. Conf. Ser. **189**, 012003(2009).
 - Abhay Ashtekar, Gen. Rel. Grav. **41**, 707(2009).
 - [17] P. Singh, J. Phys. Conf. Ser. **140**, 012005(2008).
 - [18] Jakub Mielczarek, Tomasz Stachowiak and Marek Szydlowski, Phys. Rev. D **77**, 123506(2008).
 - [19] You Ding, Yongge Ma and Jinsong Yang, Phys. Rev. Lett. **102**, 051301(2009).
 - [20] Abhay Ashtekar, Tomasz Pawlowshi and Parampreet Singh, Phys. Rev. Lett. **96**, 141301(2006).
 - [21] Abhay Ashtekar, Tomasz Pawlowshi and Parampreet Singh, Phys. Rev. D **74**, 084003(2006).
 - [22] Daris Samart and Burin Gumjudpai, Phys. Rev. D **76**, 043514(2007).
 - [23] Hao Wei and Shuang Nan Zhang, Phys. Rev. D **76**, 063005(2007).
 - [24] Xiangyun Fu, Hongwei Yu and Puxun Wu, Phys. Rev. D **78**, 063001(2008).
 - [25] Song Li and Yong-Ge Ma, arXiv: 1004.4350.
 - [26] E.J. Copeland, A.R. Liddle, and D. Wands, Phys. Rev. D **57**, 4686(1998).
 - [27] L. Amendola and S. Tsujikawa, Phys. Lett. B **660**, 125(2008).
 - [28] H.S. Zhang, H. Yu, Z.H. Zhu and Y.G. Gong, Phys. Lett. B **678**, 331(2009).
 - [29] German Olivares, Fernando Atrio-Barandela and Diego Pavon, Phys. Rev. D **77**, 063513(2008).
 - [30] German Olivares, Fernando Atrio-Barandela and Diego Pavon, Phys. Rev. D **71**, 063523(2005).
 - [31] J.H. He and B. Wang, JCAP **0806**, 010(2008).
 - [32] C. Feng, B. Wang, E. Abdalla, and R.K. Su, Phys. Lett. B **665**, 111(2008).
 - [33] D. Rowland and I. B. Whittingham, Mon. Not. R. Astron. Soc. **390**, 1719(2008).